

Lecture 11: Lattice simulations of cosmic defects

Jorge Baeza-Ballesteros and F. Torrenti

CosmoLattice school 2025 - 26th September 2025



Deutsches Elektronen
-Synchrotron Zeuthen

Overview

1. Cosmic topological defects
2. Global and local cosmic strings
 - 2.1. Cosmic strings in the continuum
 - 2.2. Lattice simulations of cosmic strings
 - 2.3. Particle and GW emission from cosmic string loops
3. Domain walls [F. Torrenti]

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3. Domain walls [F. Torrenti]

Topological defects

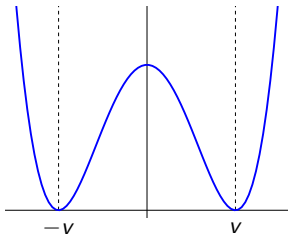
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$$S = - \int d^2x \left[\frac{1}{2} \partial_\mu \phi \partial_\mu \phi + \lambda (\phi^2 - v^2)^2 \right]$$

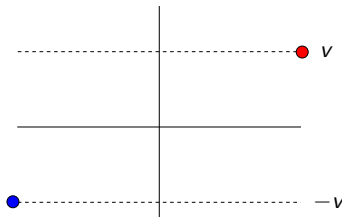
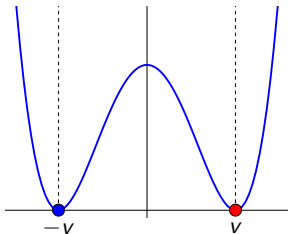


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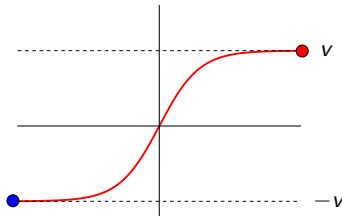
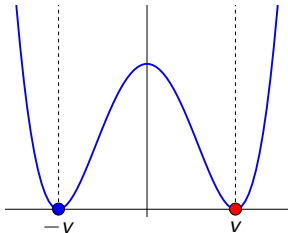


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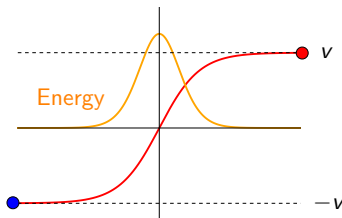
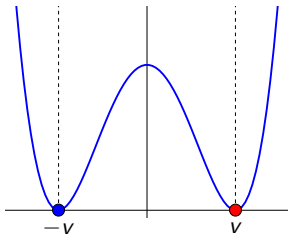
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Energy is localized
at the defects!



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► Stable due to a **conserved current** and **topological charge**

Current : $J_\mu = \varepsilon_{\mu\nu} \partial_\nu \phi$

Topological charge : $Q = \int dx J_0 = \frac{1}{2} [\phi(x = +\infty) - \phi(x = -\infty)]$

Given by **boundary values** only

Topological defects

Mathematical definition: Stable field configurations that *may* originate in field theories with a non-trivial vacuum manifold

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Same example: scalar theory in 1+1 dimensions

$$S = - \int d^2x \left[\frac{1}{2} \partial_\mu \phi \partial_\mu \phi + \frac{\lambda}{4} (\phi^2 - v^2)^2 \right]$$

Symmetry of \mathcal{L}

$$G = \mathbb{Z}_2 \quad (\phi \rightarrow -\phi)$$



Symmetry of $|0\rangle$

$$H = \mathbb{1}$$

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Vacuum manifold

$$\mathcal{M} = \{\pm v\} \cong G/H = \mathbb{Z}_2$$

Non-trivial homotopy group of \mathcal{M}

$$\pi_0(\mathcal{M}) = \mathbb{Z}_2$$

Topological defects:

Example in 3+1 dimensions: **Domain walls**

$$S_{\text{DW}} = - \int d^4x \left[\frac{1}{2} \partial_\mu \phi \partial^\mu \phi + \frac{\lambda}{4} (\phi^2 - v^2)^2 - \frac{\lambda v^4}{4} \right]$$

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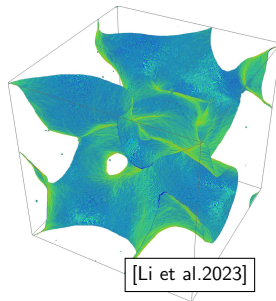
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Domain walls
- DWs are 2-dimensional surfaces embedded in 3-dimensional space
- DWs **separate inequivalent vacua**

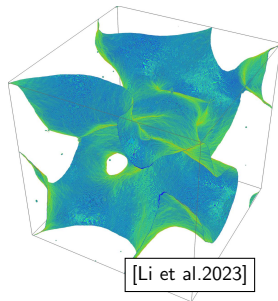


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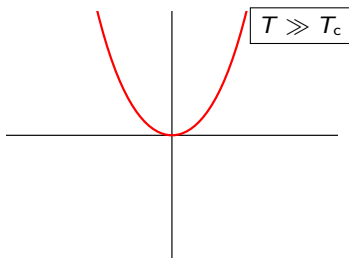


Domain walls **overclose the universe**, $\rho_{\text{DW}} \sim H \longrightarrow$ Include a **bias**!

How do domain walls form? Kibble mechanism

DWs (and other topological defects) form after spontaneous symmetry breaking

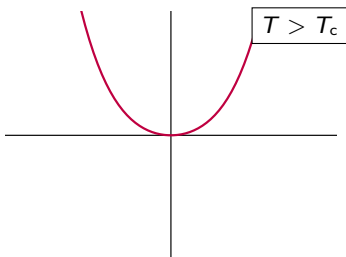
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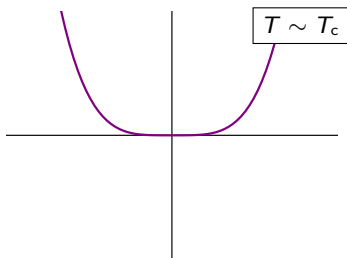
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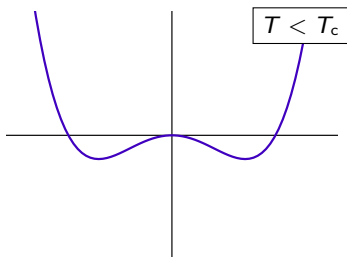
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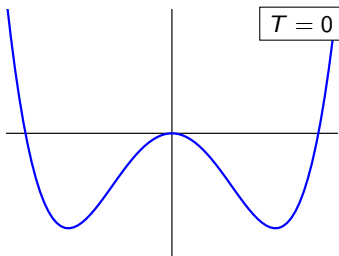
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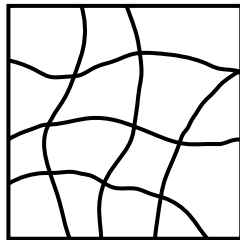
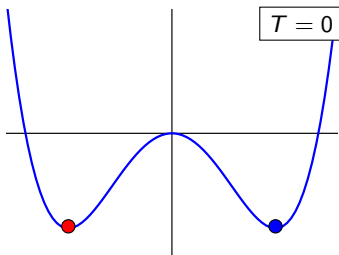
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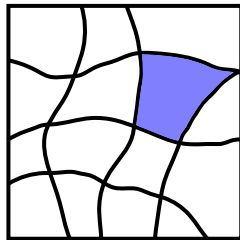
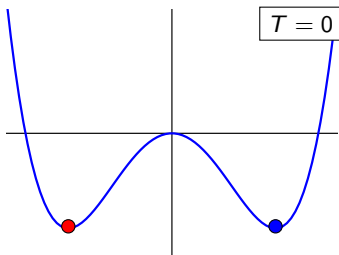
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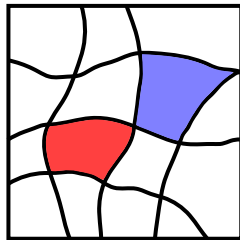
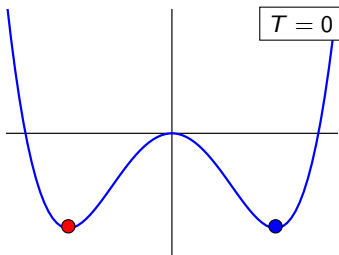
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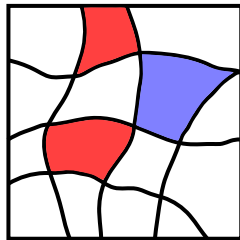
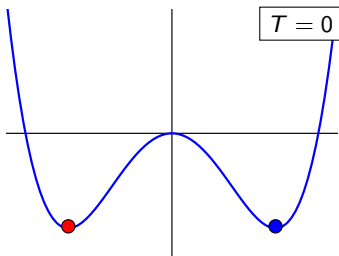
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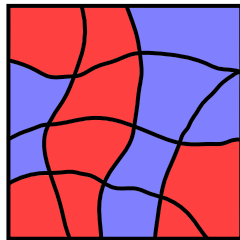
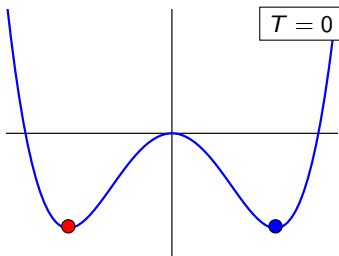
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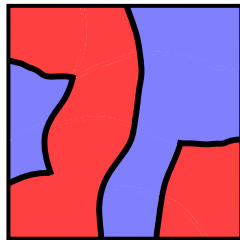
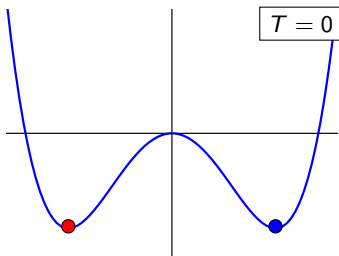


- Symmetry needs to be broken **after the end of inflation**

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Topological defects: cosmic strings

Another example:

$$S_{\text{global}} = - \int d^4x \left[(\partial_\mu \varphi)^* (\partial^\mu \varphi) + \lambda \left(|\varphi|^2 - \frac{v^2}{2} \right)^2 \right]$$

► **Symmetry breaking:**

$$\begin{aligned} G = U(1)_{\text{global}} &\longrightarrow H = \mathbb{1} \\ \mathcal{M} \cong U(1) &\longrightarrow \pi_1(\mathcal{M}) = \mathbb{Z} \end{aligned}$$

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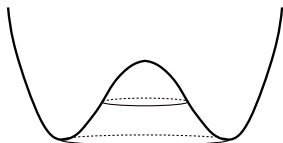
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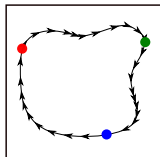
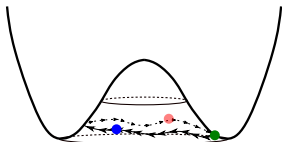
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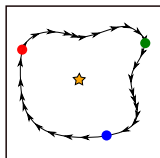
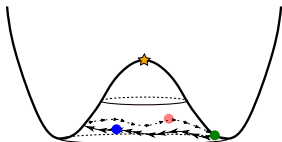
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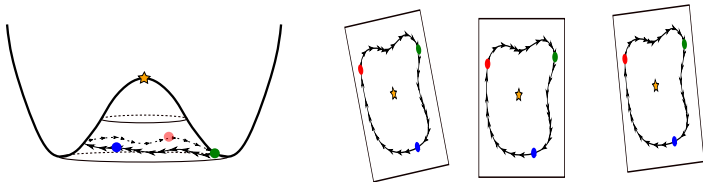
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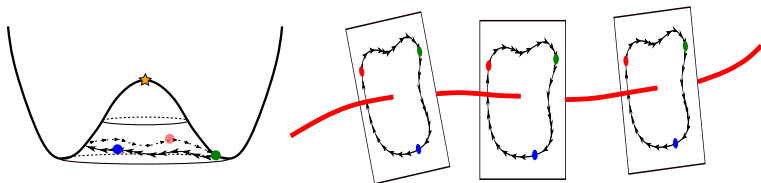
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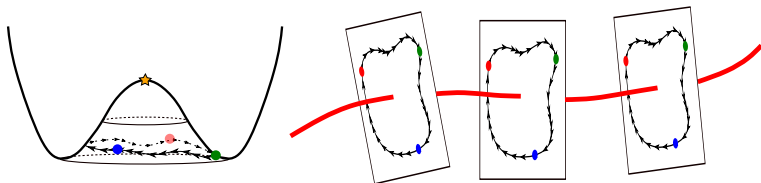
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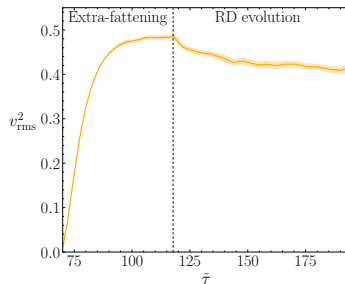
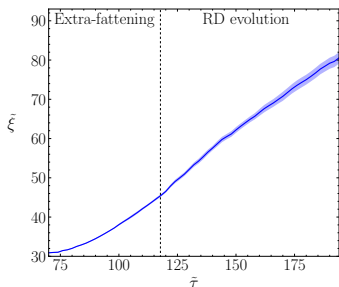


► **Many more types of strings:** **local strings**, current-carrying strings...

Scaling of topological defects

Topological defects are expected to reach a **scaling regime**:
constant density of defects per Hubble patch

- **Domain walls:** $\mathcal{A} \propto \tau$ (area parameter)
- **Cosmic strings:** $\xi \propto \tau$ (length per patch) $v_{\text{rms}}^2 \sim \text{const}$



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Global cosmic strings

Global cosmic string arise in theories with a broken **global** $U(1)$ symmetry

- **Example:** Axion theories
- Captured by a theory with a complex scalar field + mexican-hat potential

$$S_{\text{global}} = - \int d^4x \left[\partial_\mu \varphi \partial^\mu \varphi + \lambda \left(|\varphi|^2 - \frac{v^2}{2} \right)^2 \right]$$

- Equations of motion identical than any other scalar theory

$$(a^2 \varphi')' - a^2 \partial_i \partial_i \varphi = -2\lambda a^4 \left(|\varphi|^2 - \frac{v^2}{2} \right) \varphi$$

- We use **conformal time** ($\alpha = 1$)
- Both v^2 and λ can be reabsorbed

Global cosmic strings

- Contains **massive and massless particles**

Massive : $m_\chi = \sqrt{2\lambda}v \sim r_c^{-1}$

Massless : $m_\theta = 0$

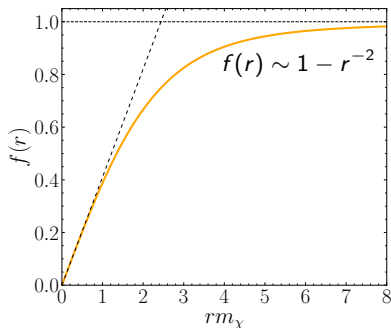
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- Analytical solution for **infinite straight string**



- Ansatz: $\varphi = f(r) \frac{v}{\sqrt{2}} e^{i\theta}$

- Logarithmic-growing tension:**

$$\mu \sim 2\pi v^2 \log\left(\frac{R}{r_c}\right)$$

- Long range interactions

Local cosmic strings

Local cosmic string arise in theory with a broken **local** U(1) symmetry

- **Example:** Grand unified theories (different than fundamental strings!)
- Captured by a theory with a complex scalar and fields + mexican-hat potential

$$S_{\text{local}} = - \int d^4x \left[(D_\mu \varphi)^* (D^\mu \varphi) + \frac{1}{4} F_{\mu\nu} F^{\mu\nu} + \lambda \left(|\varphi|^2 - \frac{v^2}{2} \right)^2 \right]$$

- Equations of motion identical than for standard U(1) theories

$$\begin{aligned} (a^2 \varphi')' - a^2 D_i D_i \varphi &= -2\lambda a^4 \left(|\varphi|^2 - \frac{v^2}{2} \right) \varphi \\ F'_{0i} - \partial_j F_{ji} &= 2a^2 e \operatorname{Im}[\varphi^* D_i \varphi] \end{aligned}$$

- Use **conformal time** ($\alpha = 1$) and **temporal gauge** ($A_0 = 0$)
- Dynamics depend on a single parameter, $\beta = e^2/2\lambda$ ($\beta = 1$, **critical case**)

Local cosmic strings

- Contains **massive particles**

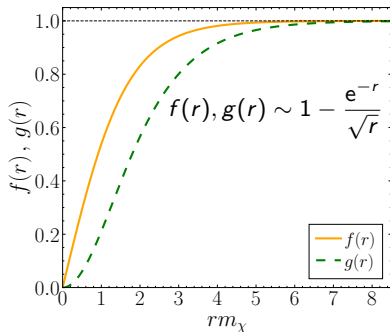
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- Ansatz:

$$\varphi = f(r) \frac{v}{\sqrt{2}} e^{i\theta}$$

$$A_i = -\frac{g(r)}{er} \hat{\theta}_i$$

- Finite tension:

$$\mu = 2\pi v^2$$

- Short range interactions

Nambu-Goto approximation

Traditionally strings have been studied in the **Nambu-Goto** limit
(**infinitely-thin string approximation**)

- **Infinite strings** only decay via the **emission of loops**
- **Loops only emit GWs** from sharp features, $P_{\text{GW}} = \Gamma G\mu$

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Accurate predictions need to incorporate
the **field-theory nature** of strings



**Lattice
simulations**

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Simulations of cosmic strings and other topological defects follow the **same basics as standard simulations of scalar and gauge theories**

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Simulations of cosmic strings and other topological defects follow the **same basics as standard simulations of scalar and gauge theories**

Some **new additions**:

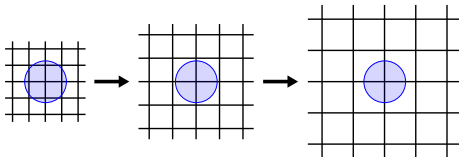
- ▶ **Resolution-preserving techniques**
- ▶ **Initial conditions** that generate a network of defects
- ▶ **Specialized observables** to keep track of the dynamics of the defects

Available in ***CosmoLattice* v2.0**

Loss of resolution on the lattice

Lattice simulations of cosmic defects suffer from **limited dynamical range**

Loss of resolution:



Simulation time

Early universe

$$R/r_c \sim \begin{array}{l} 10^{30} \text{ (axions)} \\ 10^{50} \text{ (GUTs)} \end{array}$$

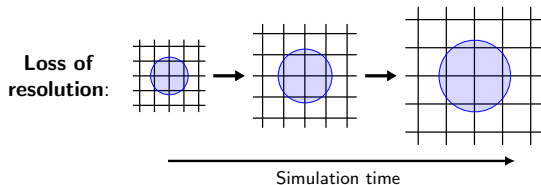


Lattice

$$R/r_c \lesssim 10^4 - 10^5$$

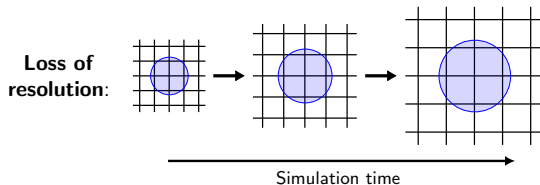
Resolution-preserving techniques

- **Fattening:** keep resolution constant

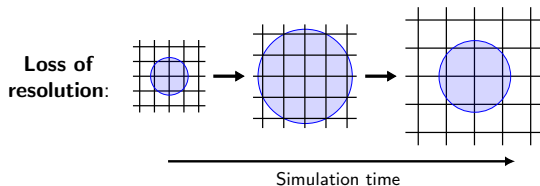


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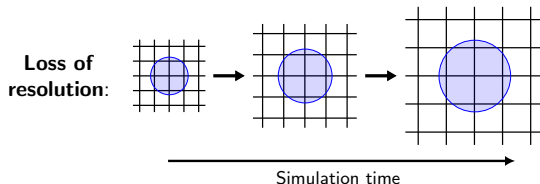


- **Extra-fattening:** grow core first, followed by physical evolution

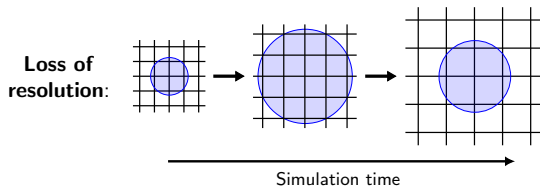


Resolution-preserving techniques

- **Fattening:** keep resolution constant



- **Extra-fattening:** grow core first, followed by physical evolution



- **Other approaches:** Adaptive mesh refinement, this is, locally increase resolution at the defect (not in *CosmoLattice* v2.0)

Resolution-preserving techniques

Generalized equations of motion:

$$\begin{aligned}(a^2 \varphi')' - a^2 D_i D_i &= -2a^{2(s+1)} \lambda \left(|\varphi|^2 - \frac{v^2}{2} \right) \varphi \\ (a^{2(1-s)} F'_{0i})' - a^{2(1-s)} \partial_j F_{ji} &= 2a^2 e \operatorname{Im} [\varphi^* D_i \varphi]\end{aligned}$$

► Different s correspond to different evolution:

- **Physical:** $s = 1 \longrightarrow w_c \sim a^{-1}$
- **Fattening:** $s = 0 \longrightarrow w_c \sim \text{const}$
- **Extra-fattening:** $s = -1 \longrightarrow w_c \sim a$

Resolution-preserving techniques

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► Equivalent to promote λ and e to time-dependent variables

$$\lambda(\tau) = \left(\frac{a}{a_0} \right)^{2(s-1)} \lambda \qquad e^2(\tau) = \left(\frac{a}{a_0} \right)^{2(s-1)} e^2$$

after redefining $A_\mu \longrightarrow A_\mu / e$

Initial conditions: phase transition vs diffusion

Physical networks live on a **scaling regime**

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- Simulate phase transition with **time-dependent potential**
 - Slow approach to scaling regime

Initial conditions: phase transition vs diffusion

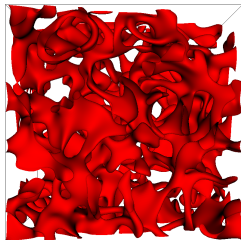
Physical networks live on a **scaling regime**

- Simulate phase transition with **time-dependent potential**
 - Slow approach to scaling regime
- Use initial phase of damping/**diffusion**:

$$\sqrt{\lambda} v \varphi' - D_i D_i \varphi = -2\lambda \left(|\varphi|^2 - \frac{v^2}{2} \right)$$

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- Empirically, **faster approach to scaling**



Initial conditions: phase transition vs diffusion

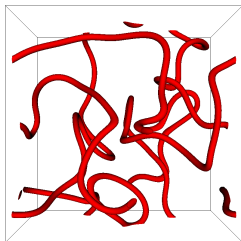
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Cosmic string observables

New observables allow to get information about the strings

Cosmic string observables

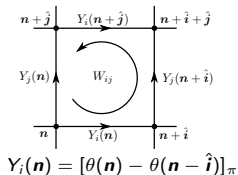
New observables allow to get information about the strings

➤ Measure the **pierced plaquettes** → Location of the **string core**

- Non-zero winding on pierced plaquettes

$$W_{ij}(\mathbf{n}) = \frac{1}{2\pi} \left[Y_i(\mathbf{n}) + Y_j(\mathbf{n} + \hat{i}) - Y_i(\mathbf{n} + \hat{j}) - Y_j(\mathbf{n}) \right]$$

- **Total length** of strings: $L_w = \frac{2}{3} \delta x \#_{\text{pierced}}$
Manhattan effect



Cosmic string observables

New observables allow to get information about the strings

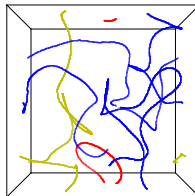
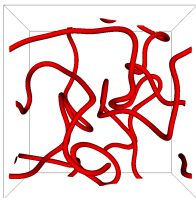
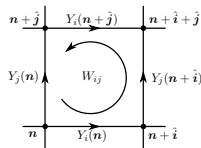
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- **Total length** of strings: $L_w = \frac{2}{3} \delta x \#_{\text{pierced}}$

- Real-time quantitative information of network structure



Cosmic string observables

New observables allow to get information about the strings

- String energy components using **weight functions**

$$E_{\text{str}} = \int d^3x W[\varphi] \rho(\mathbf{x}) \quad \text{with} \quad W[\varphi] \propto V[\varphi] \Theta\left(|\varphi|^2 - \frac{1}{2}\right)$$

- **Estimates of total length and mean-squared velocity**

$$L_{\text{str}} = \frac{1}{\mu} \frac{E_{\text{str}} + f_V L_{\text{str}}}{1 - f_V}$$

Global string :

$$v_{\text{rms}}^2 = \frac{E_{\text{str}} + L_{\text{str}}}{E_{\text{str}} + f_V L_{\text{str}}}$$

where μ and f_V are computed from the infinite string solution and

$$L_{\text{str}} = K_{\text{str}} - G_{\text{str}} - V_{\text{str}}$$

Overview

1. Cosmic topological defects

2. Global and local cosmic strings

2.1. Cosmic strings in the continuum

2.2. Lattice simulations of cosmic strings

2.3. Particle and GW emission from cosmic string loops

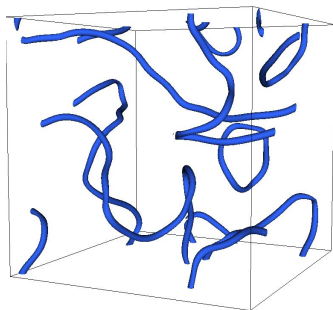
3. Domain walls [F. Torrenti]

Decay of isolated loops

We have studied the **decay of loops into particles and GWs**

[JBB, Copeland, Figueroa and Lizzarraga, 2023 and 2024]

Local string



- ▶ We studied **isolated global and local loops in flat background** to avoid loss of resolution

- ▶ Analyze emission of **particles and GWs** simultaneously

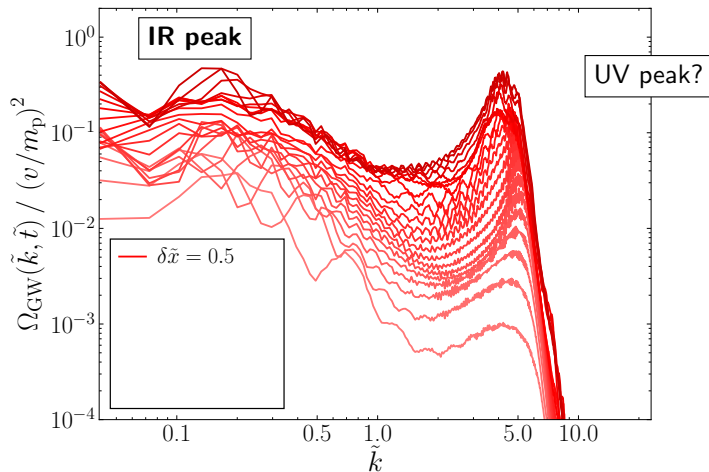
- ▶ Reached separation of scales

$$L_w/r_c \sim 10^4$$

- ▶ **Goal:** combine results for the GW emission with history of loop number density to predict GWB

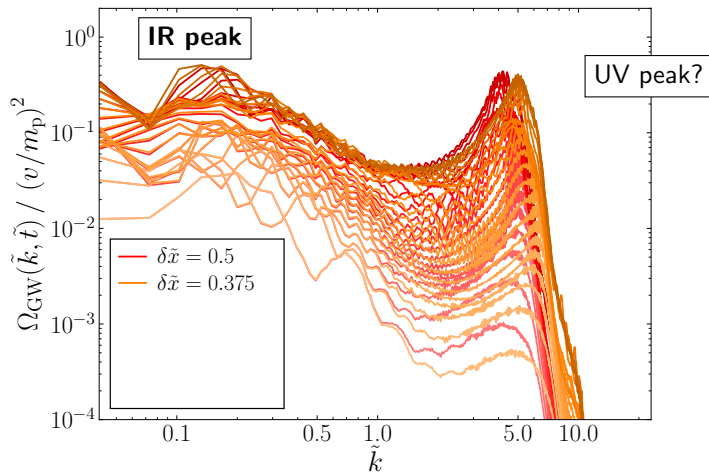
GW emission from loops: UV resolution

Global loop: Fixed loop and lattice size + vary δx



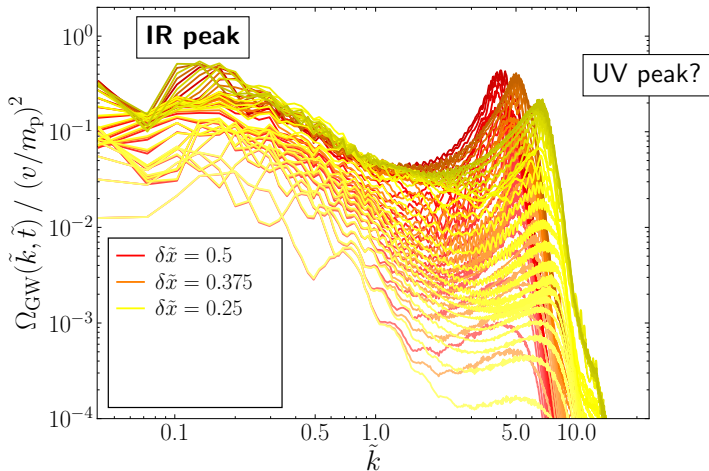
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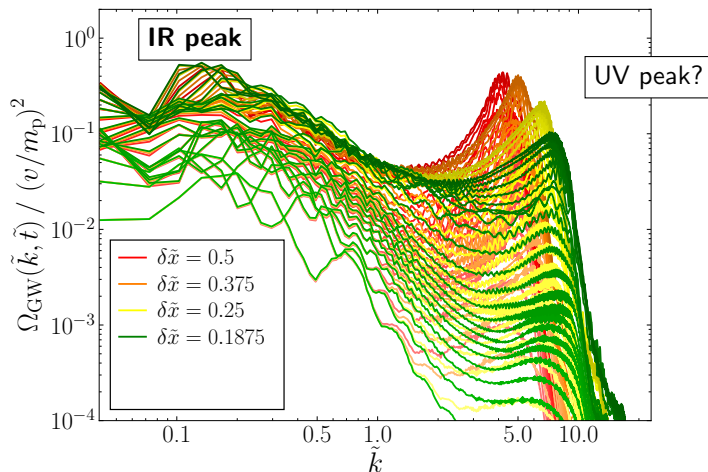
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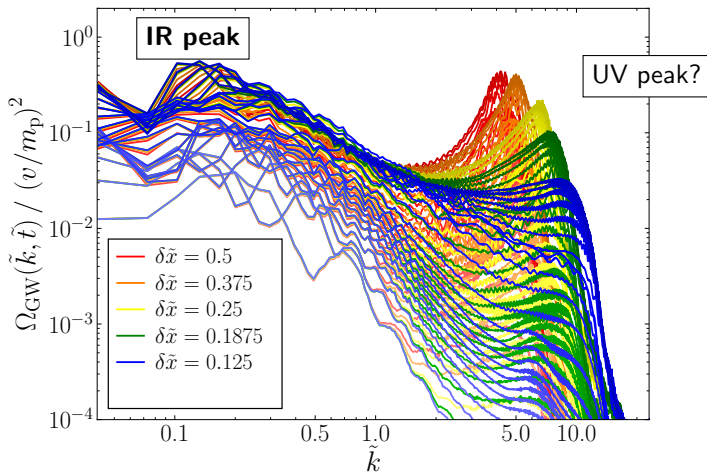
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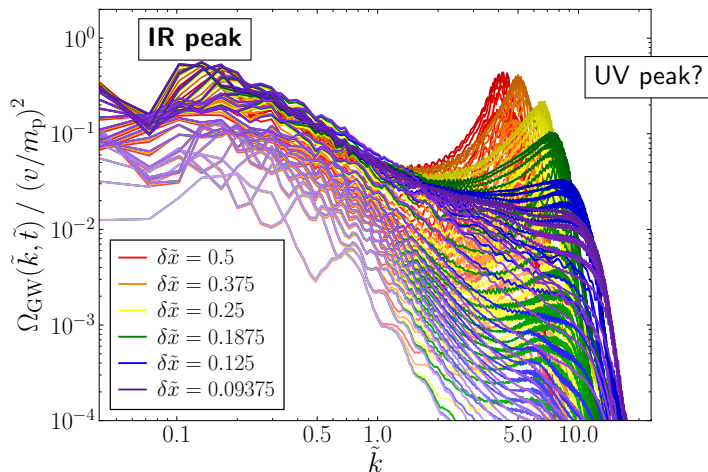
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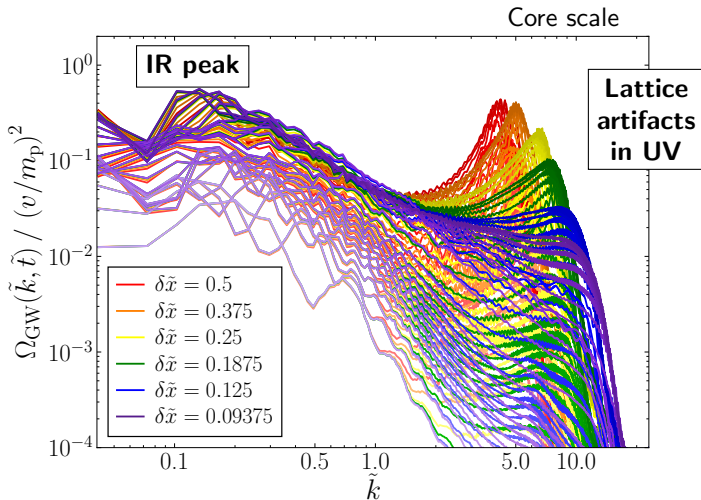
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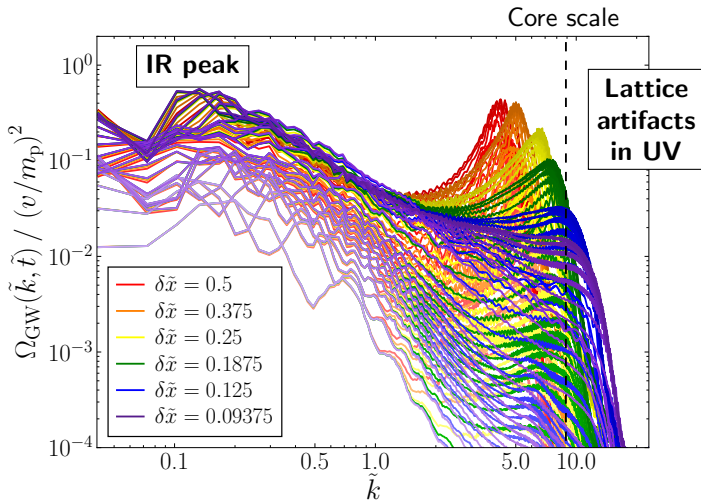
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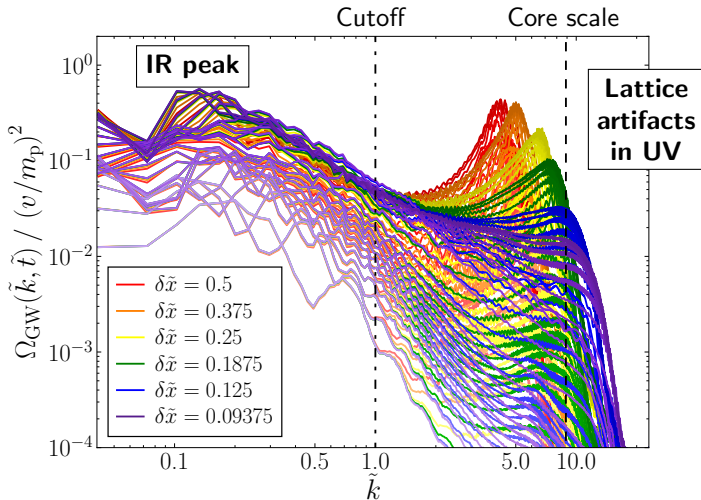
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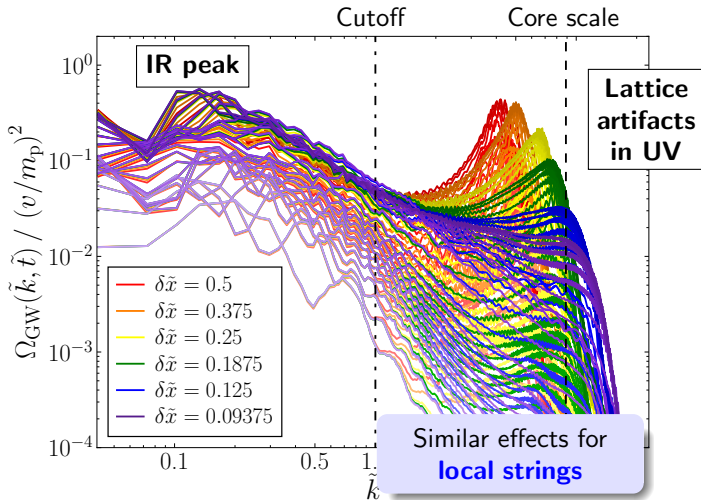
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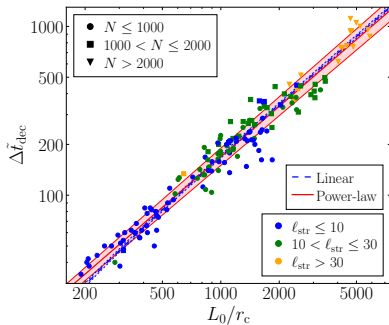
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Particle and GW emission from local loops

We have simulated 100+ **local loops**, reaching $L_0/r_c \lesssim 7000$ [JBB et al. 2024]

Particle emission

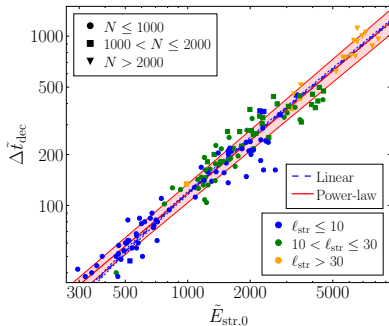


$$\Delta \tilde{t}_{\text{dec}} \propto \tilde{L}_0$$

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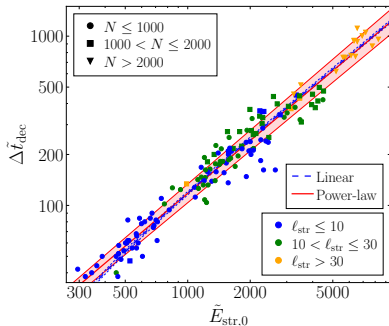


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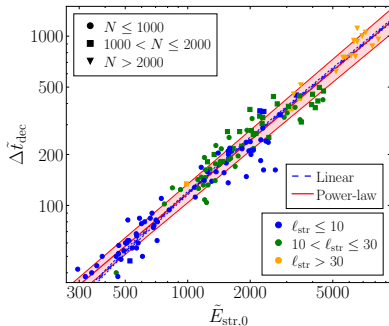


$$P_{\text{part}} = 7.70(0.16)v^2$$

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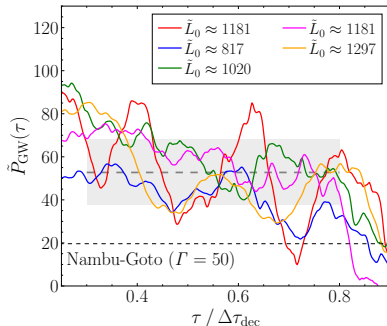
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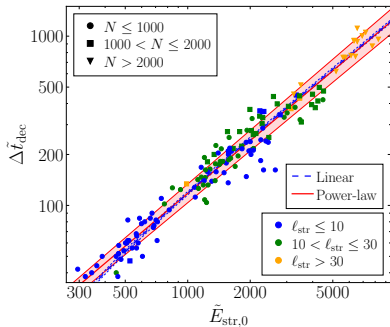


$$P_{\text{GW}} = 54(16) \times v^2 \left(\frac{v}{m_p} \right)^2$$

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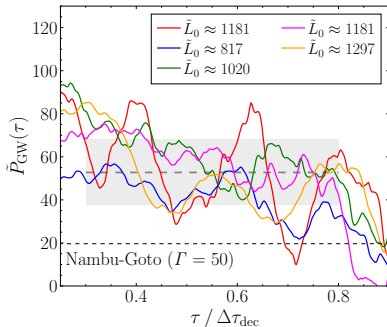
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$$P_{\text{GW}} = 54(16) \times v^2 \left(\frac{v}{m_p} \right)^2$$

GW emission is **very suppressed** compared to NG predictions

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